

Status and prospects of non-Riemannian cosmology

Dirk Puetzfeld
(Iowa State University)

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Why cosmology?

- Local tests in rather good compliance with GR (maybe not on very small scales)
- New effects likely to show up on **large scales** and/or **high energies**
- Observations force us to introduce concepts like **dark matter** and **dark energy** within the standard cosmological model
- So far **no** direct detection of a dark matter particle
- So far **no** convincing theoretical explanation for the dark energy component

Alternative approaches to cosmology

- Scalar fields

Ratra & Peebles 1988, Wetterich 1988

- Time varying cosmological constant

Oezer & Taha 1987, Vishwakarma 2001

- K-essence

Chiba et al. 2001, Armendariz-Picon et al. 2000

- Phantom energy

Caldwell 2002

- Chaplygin gas

Barrow 1990, Hassaine et al. 2001, Fabris et al. 2002

- Cardassian expansion

Freese & Lewis 2002, Zhu & Fujimoto 2002

- Brane world models

Randall & Sundrum 1999, Deffayet et al. 1999

- Non-symmetric gravity

Moffat 1997, 2001, 2004

- **Non-Riemannian models**

Why go beyond Riemannian gravity?

- Curved spacetime (GR) \sim dynamics of masspoints and light
- Intrinsic properties of particles suggest couplings to new fields
- Simple analogy:



why not more complex?

Why non-Riemannian gravity ?

- Elementary particles may be classified by the Poincaré group by their **mass** and **spin**
- **Mass** connected with the **translational** part of the Poincaré group
- **Spin** connected with the **rotational** part of the Poincaré group
- Mass and spin are **elementary notions**, each with an analogous standing not reducible to that of the other
- Distributing **mass-energy** and **spin** over space-time leads to the field theoretical notions of an **energy-momentum** tensor and **spin angular momentum** tensor of matter
- In macrophysical limit, mass or **energy-momentum** adds up because of its **monopole** character
- **Intrinsic spin** has **dipole** character and usually averages out in the macroscopic limit (this is one of the reasons why GR is successful on macroscopic scales)
- In analogy to the coupling of the energy-momentum to the metric one expects that that also spin angular momentum couples to a new quality which is linked to the geometry of spacetime



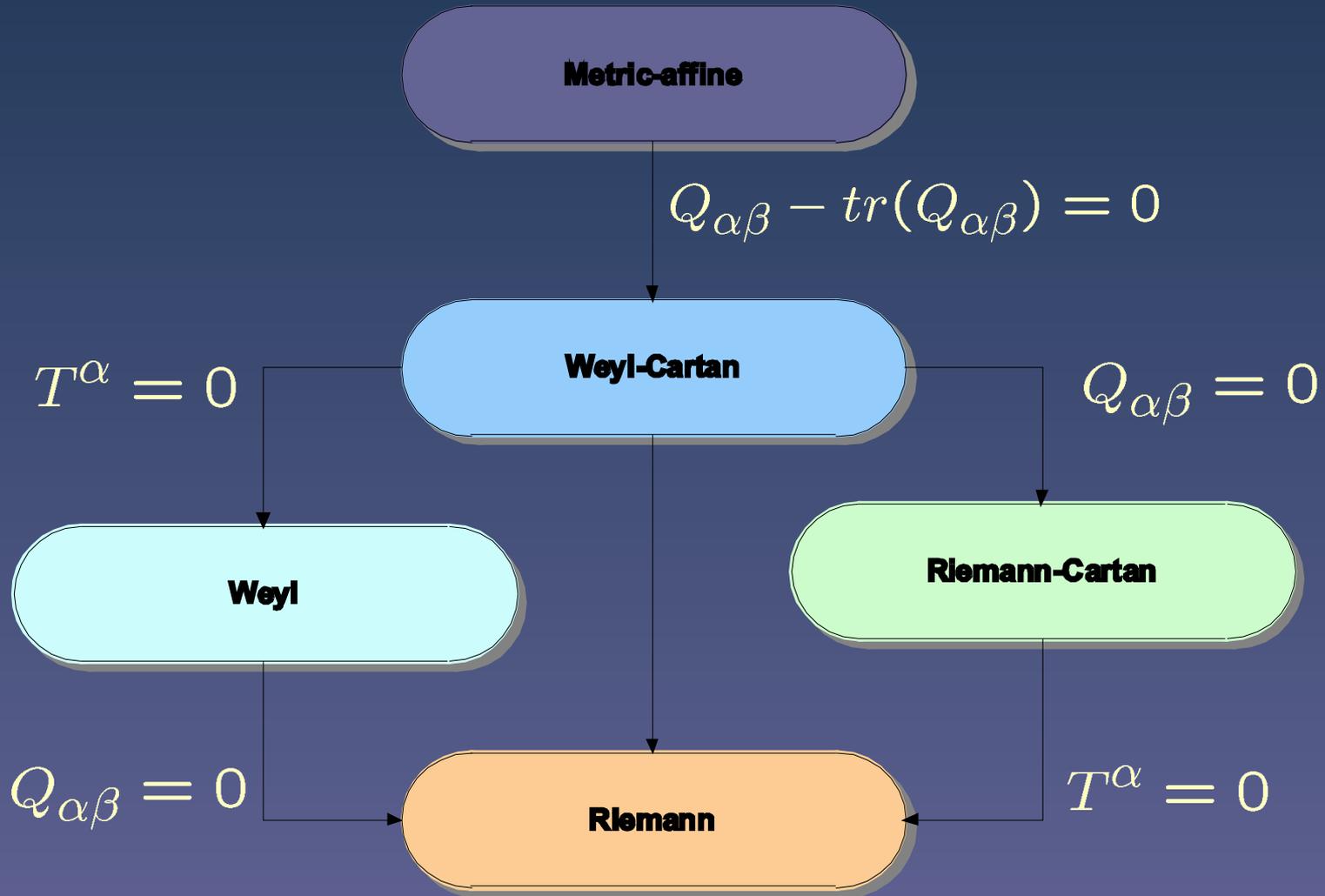
Riemann-Cartan spacetime

Metric-affine gravity (MAG)

Potentials	Field strenghts	Gauge currents	Excitations	Matter currents
$g_{\alpha\beta}$	$Q_{\alpha\beta}$	$m^{\alpha\beta}$	$M^{\alpha\beta}$	$\sigma^{\alpha\beta}$
ϑ^α	T^α	E_α	H_α	Σ_α
$\Gamma_\alpha{}^\beta$	$R_\alpha{}^\beta$	$E^\alpha{}_\beta$	$H^\alpha{}_\beta$	$\Delta^\alpha{}_\beta$

$$L = V_{\text{MAG}}(g_{\alpha\beta}, \vartheta^\alpha, Q_{\alpha\beta}, T^\alpha, R_\alpha{}^\beta) + L_{\text{mat}}(g_{\alpha\beta}, \vartheta^\alpha, \psi, D\psi)$$

Spacetime types



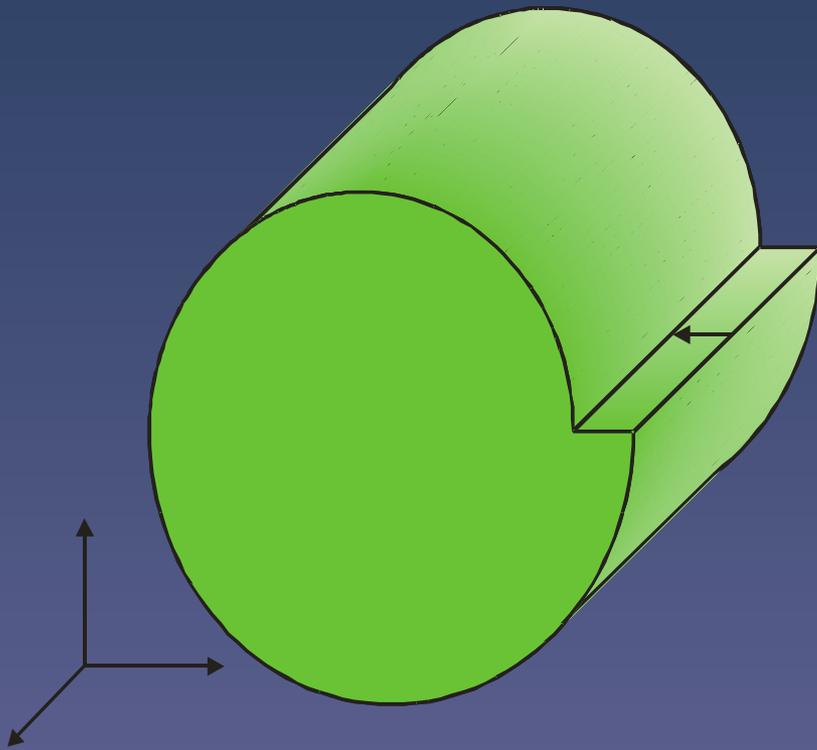
The general MAG connection

$$\Gamma_{\alpha\beta} = \tilde{\Gamma}_{\alpha\beta} + N_{\alpha\beta} =$$

$$\underbrace{\frac{1}{2}dg_{\alpha\beta} + (e_{[\alpha]}dg_{\alpha]\gamma})\vartheta^\gamma + e_{[\alpha]}C_{\beta]} - \frac{1}{2}(e_\alpha]e_\beta]C_\gamma)\vartheta^\gamma}_{\text{Levi-Civita connection}}$$

$$\underbrace{-e_{[\alpha]}T_{\beta]} + \frac{1}{2}(e_\alpha]e_\beta]T_\gamma)\vartheta^\gamma + \frac{1}{2}Q_{\alpha\beta} + (e_{[\alpha]}Q_{\beta]\gamma})\vartheta^\gamma}_{\text{Distortion}}$$

Analogy with elasticity theory and theory of defects



Torsion & Nonmetricity \sim Dislocations

$$b^i \sim \int \int dx^k \wedge dx^l T^i$$

In some cases torsion can be interpreted as the surface density of the Burgers vector, i.e. it is proportional to the dislocation density of an elastic medium

Kröner 1958, Hehl 1967

Field equations

$$\frac{\delta L_{\text{mat}}}{\delta \psi} = 0$$

$$\mathbf{0} \quad DM^{\alpha\beta} - m^{\alpha\beta} = \sigma^{\alpha\beta}$$

$$\mathbf{I} \quad DH_{\alpha} - E_{\alpha} = \Sigma_{\alpha}$$

$$\mathbf{II} \quad DH^{\alpha}_{\beta} - E^{\alpha}_{\beta} = \Delta^{\alpha}_{\beta}$$

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The general MAG Lagrangian

$$\begin{aligned}
 V_{\text{MAG}} = & \frac{1}{2\kappa} [-a_0 R^{\alpha\beta} \wedge \eta_{\alpha\beta} - 2\lambda\eta + T^\alpha \wedge \star \left(\sum_{I=1}^3 a_I {}^{(I)}T_\alpha \right) \\
 & + Q_{\alpha\beta} \wedge \star \left(\sum_{I=1}^4 b_I {}^{(I)}Q^{\alpha\beta} \right) \\
 & + b_5 \left({}^{(3)}Q_{\alpha\gamma} \wedge \vartheta^\alpha \right) \wedge \star \left({}^{(4)}Q^{\beta\gamma} \wedge \vartheta_\beta \right) \\
 & + 2 \left(\sum_{I=2}^4 c_I {}^{(I)}Q_{\alpha\beta} \right) \wedge \vartheta^\alpha \wedge \star T^\beta] \\
 -\frac{1}{2\rho} R^{\alpha\beta} \wedge \star [& \sum_{I=1}^6 w_I {}^{(I)}W_{\alpha\beta} + \sum_{I=1}^5 z_I {}^{(I)}Z_{\alpha\beta} + w_7 \vartheta_\alpha \wedge (e_\gamma] {}^{(5)}W^{\gamma}_\beta) \\
 & + z_6 \vartheta_\gamma \wedge (e_\alpha] {}^{(2)}Z^{\gamma}_\beta) + \sum_{I=7}^9 z_I \vartheta_\alpha \wedge (e_\gamma] {}^{(I-4)}Z^{\gamma}_\beta)]
 \end{aligned}$$

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 & + 2 \left(\sum_{I=2}^4 c_I {}^{(I)}Q_{\alpha\beta} \right) \wedge \vartheta^\alpha \wedge \star T^\beta \\
 & - \frac{1}{2\rho} R^{\alpha\beta} \wedge \star \left[\sum_{I=1}^6 w_I {}^{(I)}W_{\alpha\beta} + \sum_{I=1}^5 z_I {}^{(I)}Z_{\alpha\beta} + w_7 \vartheta_\alpha \wedge (e_\gamma] {}^{(5)}W^{\gamma\beta} \right. \\
 & \left. + z_6 \vartheta_\gamma \wedge (e_\alpha] {}^{(2)}Z^{\gamma\beta} \right) + \sum_{I=7}^9 z_I \vartheta_\alpha \wedge (e_\gamma] {}^{(I-4)}Z^{\gamma\beta} \left. \right]
 \end{aligned}$$

Timeline NRC models 1972-1985



Timeline NRC models 1986-1999



Timeline NRC models 2001-2004

Models enter quantitative regime



First SNIa parameter estimates in WC model

NGT model

Extended WC model

			Szydlowski
	Minkevich		Böhmer
	Puetzfeld	Minkevich	Scholz
	Shapiro	Vereshchagin	Moffat
Moffat	Capozziello	Babourova et al.	Miritzis
Puetzfeld et al.	Puetzfeld	Capozziello et al.	Puetzfeld et al.

2001

2002

2003

2004



Summary NRC

- Field equations manageable for cosmological models which make use of a constrained Lagrangian
- Possible to construct viable models (passed several cosmological tests: SN Ia, X-ray cluster, FR IIb radio galaxies, BBN)
- Interesting correspondences between NRC, anisotropic, and braneworld models
- So far **no** solution of the dark energy/matter problem

Outlook NRC

Near future

- Sophisticated fluid models offer a good framework for cosmology
- Incorporate the full hyperfluid symmetries

—————→ **Continue search for genuine effects**

Far future

- Investigate full MAG Lagrangian (difficult!)
- Work out perturbed field equations

“...the question whether this [spacetime] continuum is Euclidean or structured according to the Riemannian scheme or **still otherwise** is a genuine physical question which has to be answered by experience rather than being a mere convention to be chosen on the basis of expediency.”

A. Einstein, *Geometrie und Erfahrung*, translation by F.W. Hehl

„From a fundamental viewpoint it is totally wrong to aim at basing a theory only on observable quantities. For in reality it is just the other way around. Only the theory decides about what can be observed.“

„Aber vom prinzipiellen Standpunkt aus ist es ganz falsch, eine Theorie nur auf beobachtbaren Größen gründen zu wollen. Denn es ist ja in Wirklichkeit genau umgekehrt. Erst die Theorie entscheidet darüber, was man beobachten kann.“

Einstein according to Heisenberg (1979)

“It seems to me...that it is not so much the linearity or non-linearity which forms the heart of the matter, but the very fact that here a **more general group** than the Lorentz group is present....”

W.Pauli, *Helv. Phys. Acta Suppl.* 4 (1955) 261-267, translation by C. Kiefer